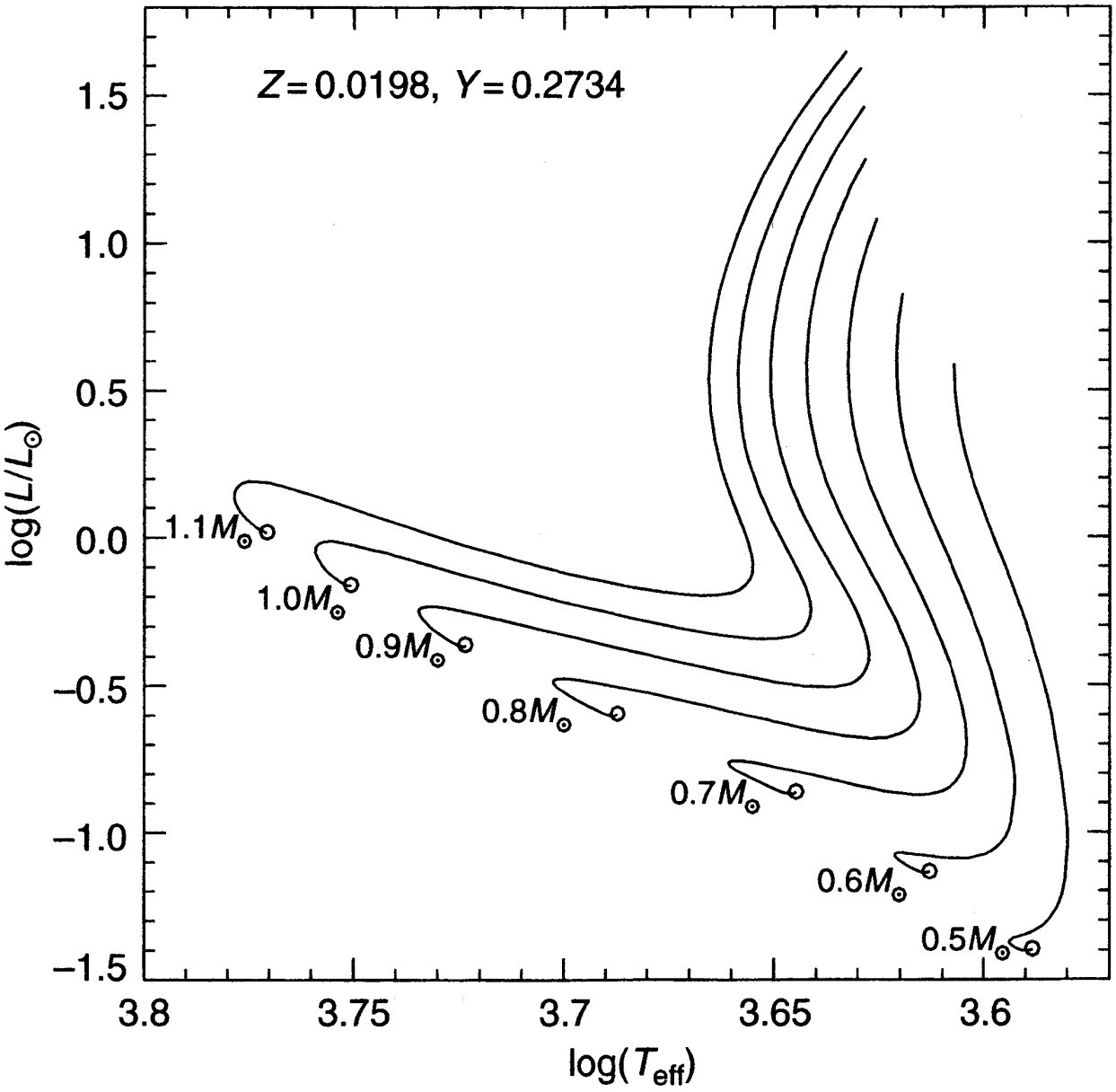

Salaris & Cassisi, Evolution of Stars and Stellar Populations, Wiley, 2005

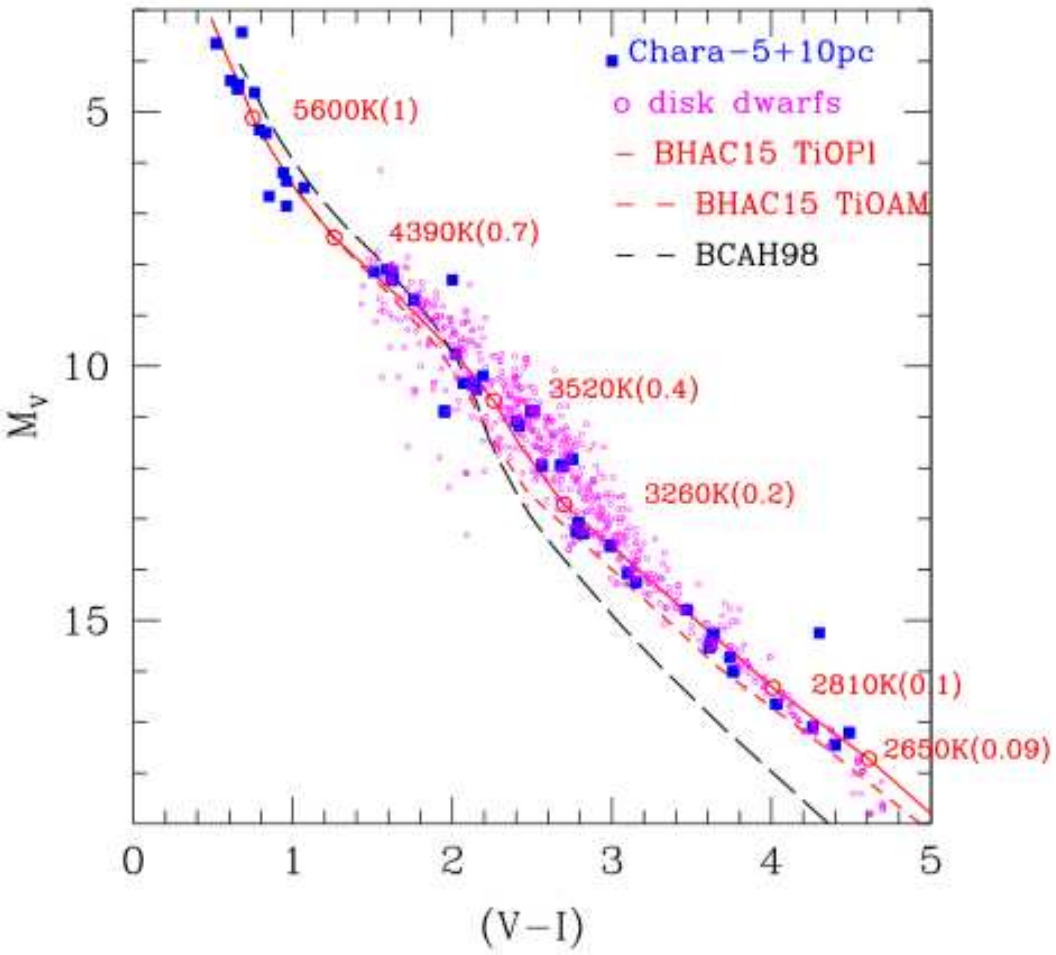

Fig. 4. Comparison of 1 Gyr isochrones with disk dwarfs in an optical CMD. Present models based on the Plez TiO (solid red) and the AMES TiO (dash red) linelists are shown, along with the BCAH98 models (long dash black). The magenta open circles are disk objects with parallax (Monet et al., 1992; Leggett, 1992) and the blue squares are from the recent 5 pc and 10 pc solar neighbourhood samples of Cantrell et al. (2013) with more accurate parallaxes and cleaned of binary systems. The numbers close to the red open circles on the solid curve give $T_{\rm eff}$ and mass (in $M_{\odot}$ in the brackets) for selected models.

Fig. 1.— Baraffe, Homeier, Allard & Chabrier, 2015, Fig. 4

Fig. 2.— The "jewel-box" nebula, NGC 3603, a region of intense star formation, as viewed by HST (Hubble Heritage image).

Fig. 3.— The "jewel-box" nebula, NGC 3603, a region of intense star formation, as viewed by HST (Hubble Heritage image).

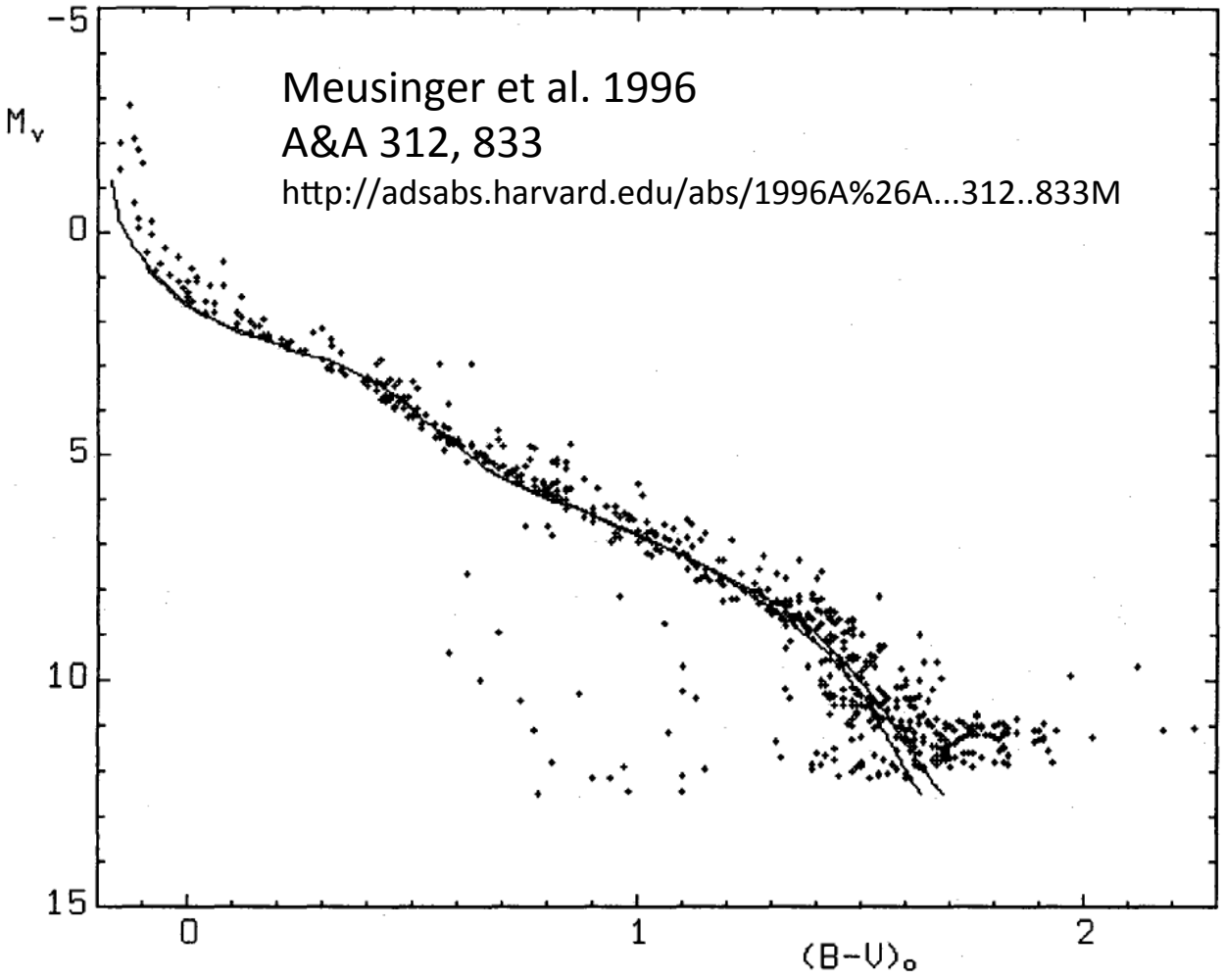

Fig. 3. The CM diagram for 627 stars in our basic Pleiades sample from Sect. 3.1 (not shown is the white dwarf LB-1497 (Luyten & Herbig 1960; Schilbach et al. 1995a) with $M_v = 11.09$, $(B-V)_0 = -0.42$). The curves are the isochrones from the models no. 27 (upper curve) and 33 (lower curve), respectively (see Sect. 4).

Meusinger et al. 1996
A&A 312, 833
http://adsabs.harvard.edu/abs/1996A%26A...312..833M

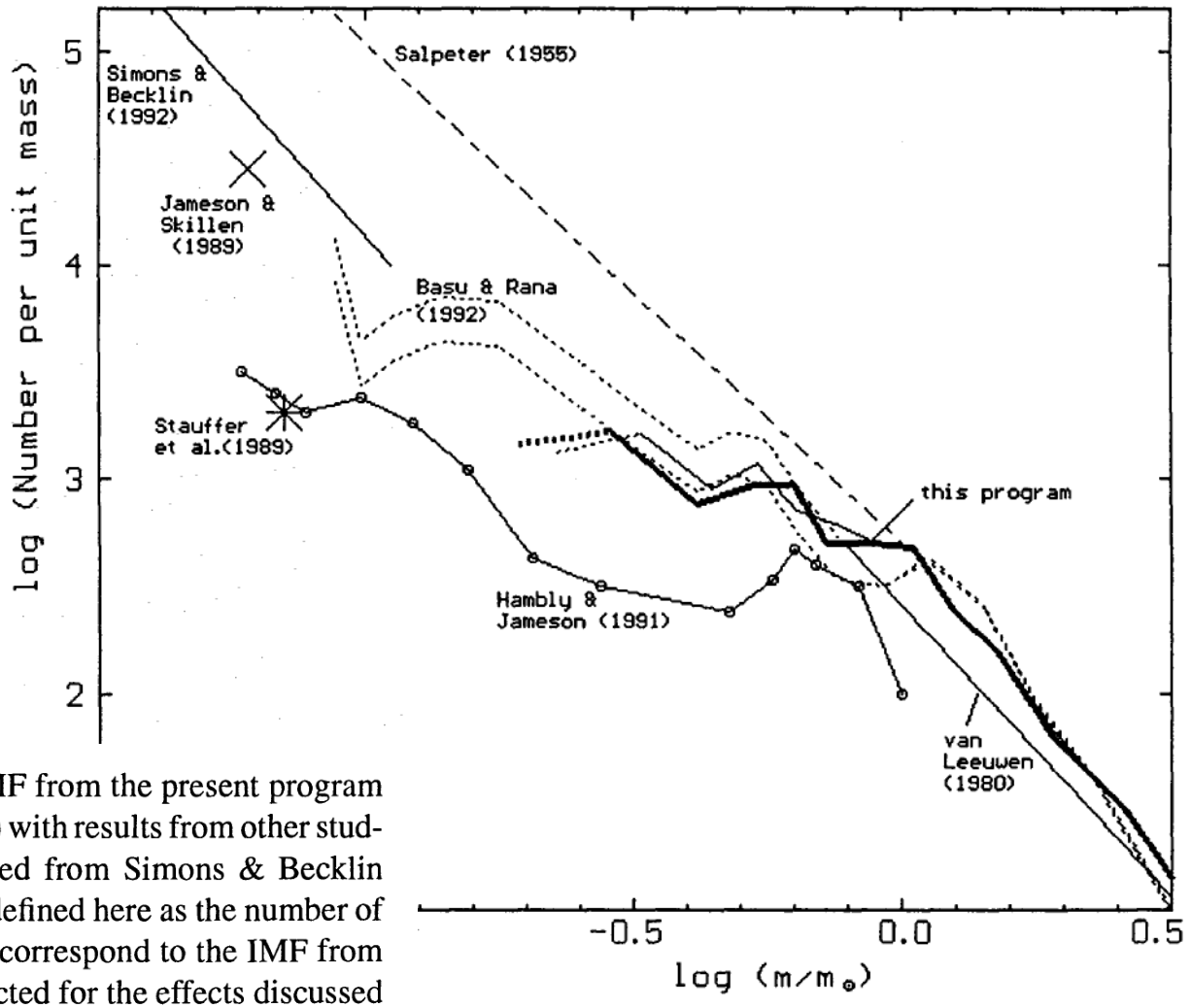

Fig. 11. A comparison of the Pleiades IMF from the present program (the weighted-averaged IMF from Fig. 10) with results from other studies. The style of representation is adopted from Simons & Becklin (1992, their Fig. 8). Note that the IMF is defined here as the number of stars per mass interval. The three curves correspond to the IMF from the uncorrected LF (I), from the LF corrected for the effects discussed in Sect. 3.4, but for the incomplete coverage of the Pleiades field (II), and from the corrected final LF (III). The dotted parts of the curves suffer from incompleteness near the limiting magnitude of the survey. The dashed line and the dotted lines are representations of the field star IMF. (The upper dotted curve corresponds to an age of 10 Gyr for the Galactic disk, the lower one to 15 Gyr. These curves are normalized in such a way that they match the Pleiades IMF at $\log m \approx 0$.)

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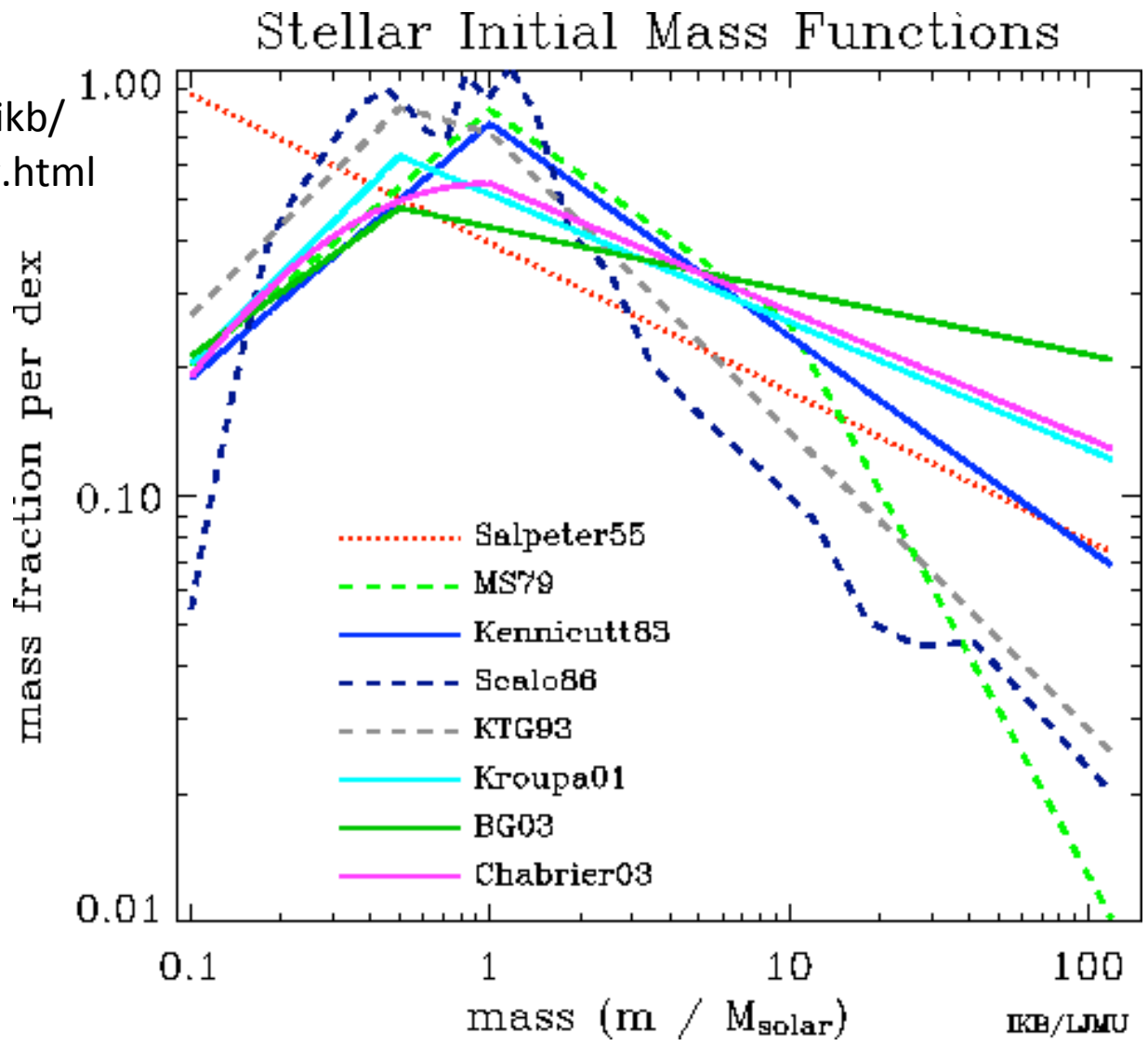
http://www.astro.ljmu.ac.uk/~ikb/research/imf-use-in-cosmology.html



This figure compares the IMFs by ploting mass fraction per dex versus mass, i.e., normalized so that the integral under each curve is unity.
They are assumed to be valid from 0.1 to 120 solar masses.