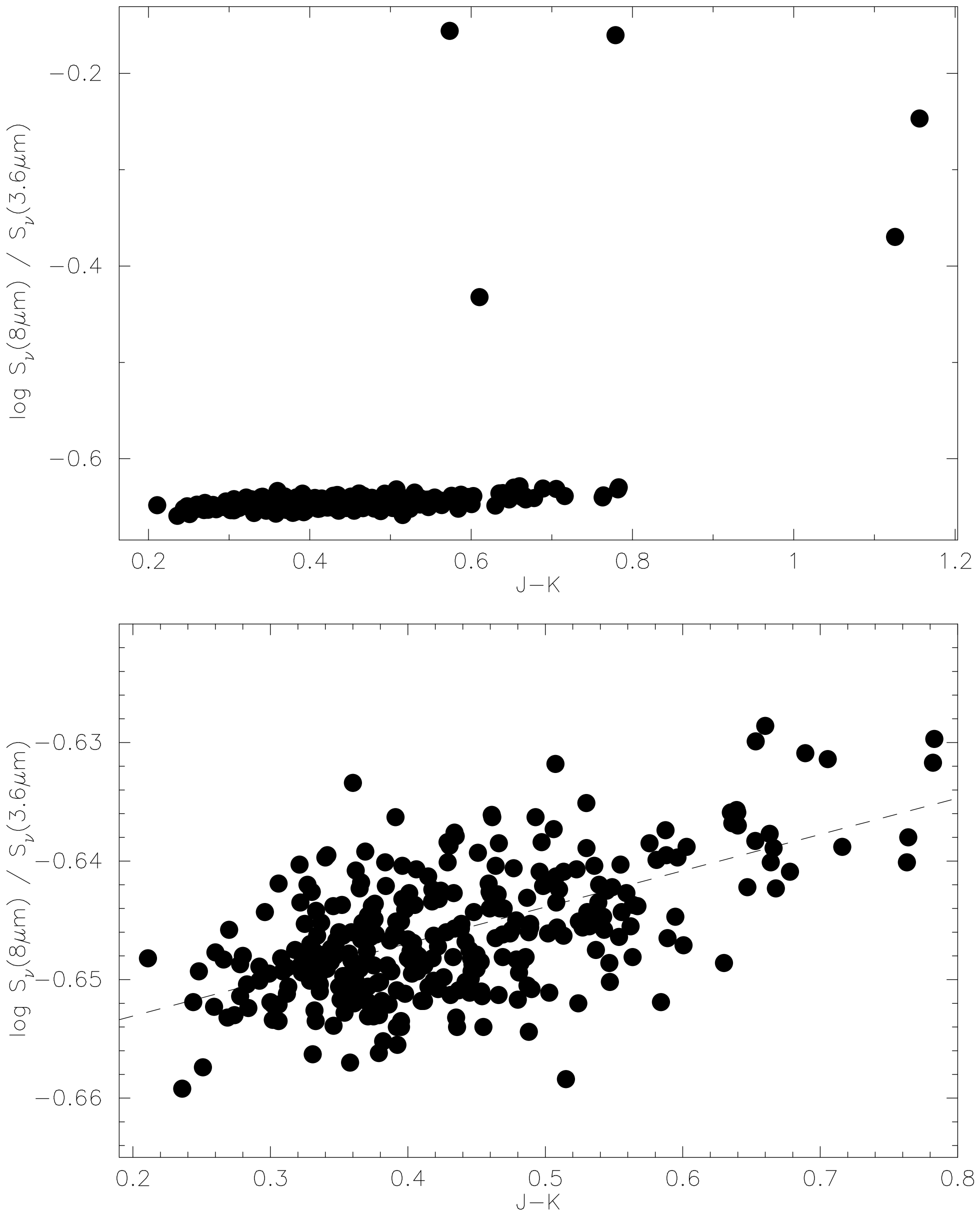

Fig. 1.— Dereddened color-color diagrams showing the $J - K_s$ versus the 8μm to 3.6μm flux ratio. The top panel shows data for all 314 stars in the sample, and the bottom panel excludes the five stars with flux ratio greater than 0.25 that have infrared excesses. The dashed lin the bottom panel represents a linear fit the data.

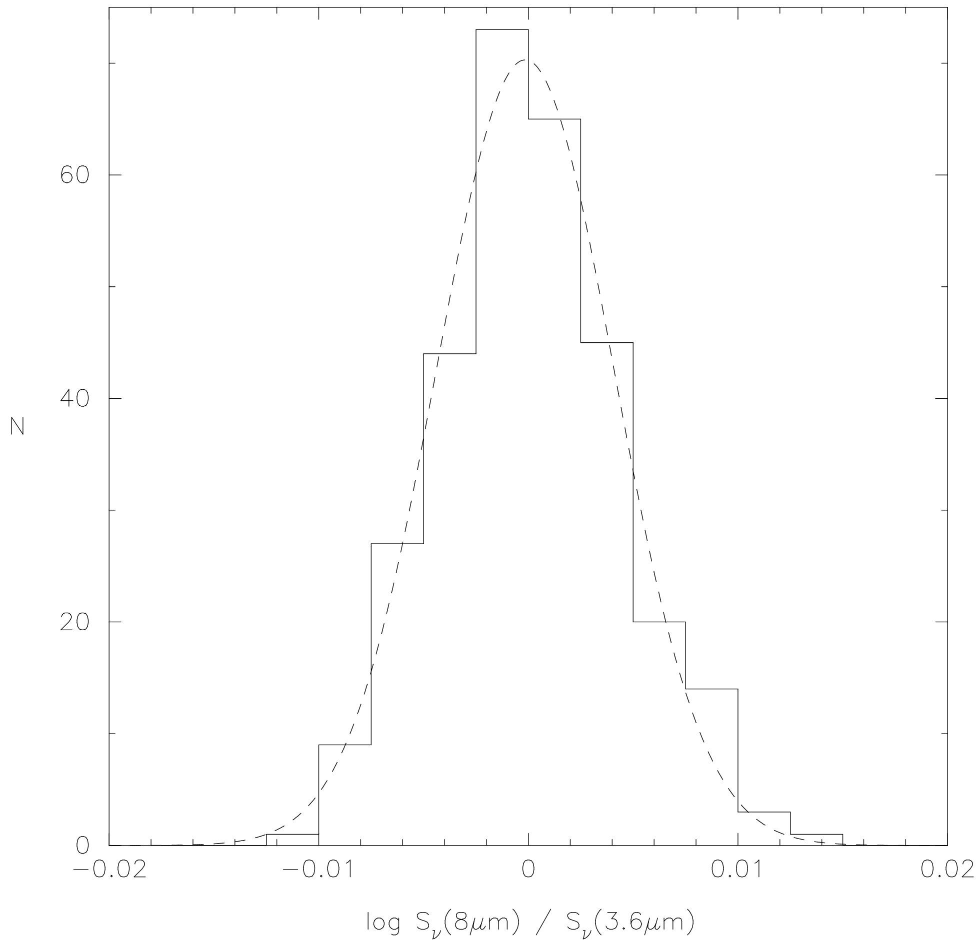

Fig. 2.— Histogram of the residual 8μm to 3.6μm flux ratio after removing the linear fit shown in the bottom panel of Fig. 1. Five stars with clear 8μm excesses are off scale. The dashed curve shows the best-fit gaussian with a dispersion of 0.0044 dex.

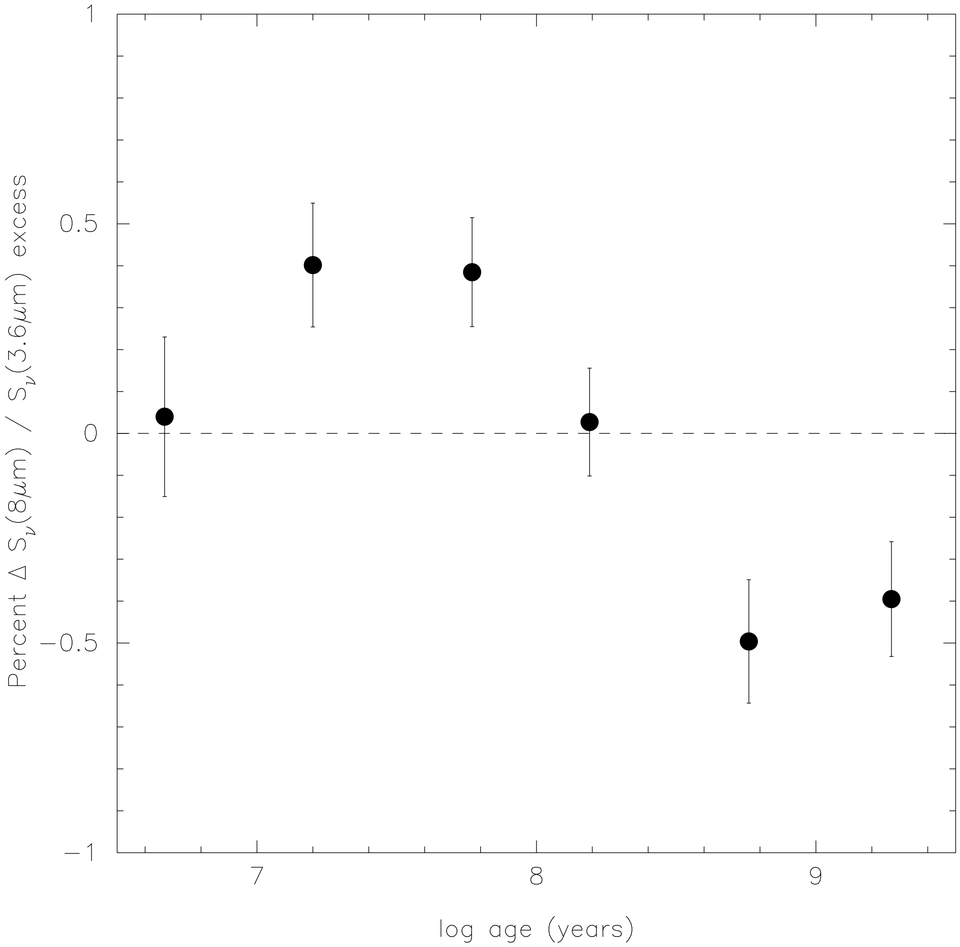

Fig. 3.— Average R_8 percent excess as a function of stellar age relative to the photosphere established by the dashed line in Figure 1 after excluding five optically thick disk. The results were used to place a 3σ upper limit of 1.3% to the 8μm excess around the average star young than 100 Myr.

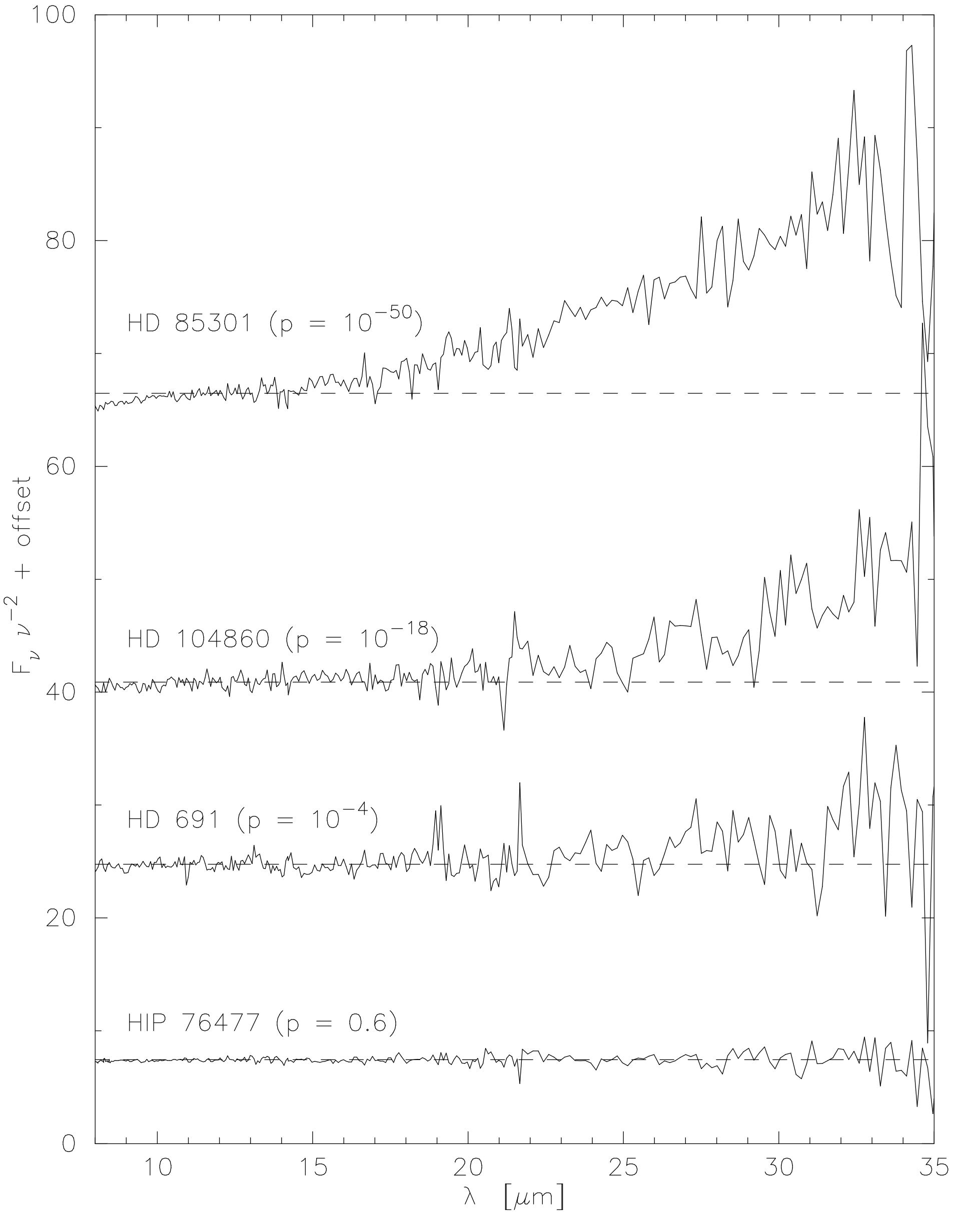

Fig. 4.— Example observed IRS low resolution spectra for four stars. The dashed line indicates the $F_\nu \nu^{-2}$ = constant spectrum normalized between 10 and 14μm. The probability, p, that a Kurucz photosphere can fit the observed spectra is indicated. Low values of p indicate the greater probability that the source has emission from warm dust.

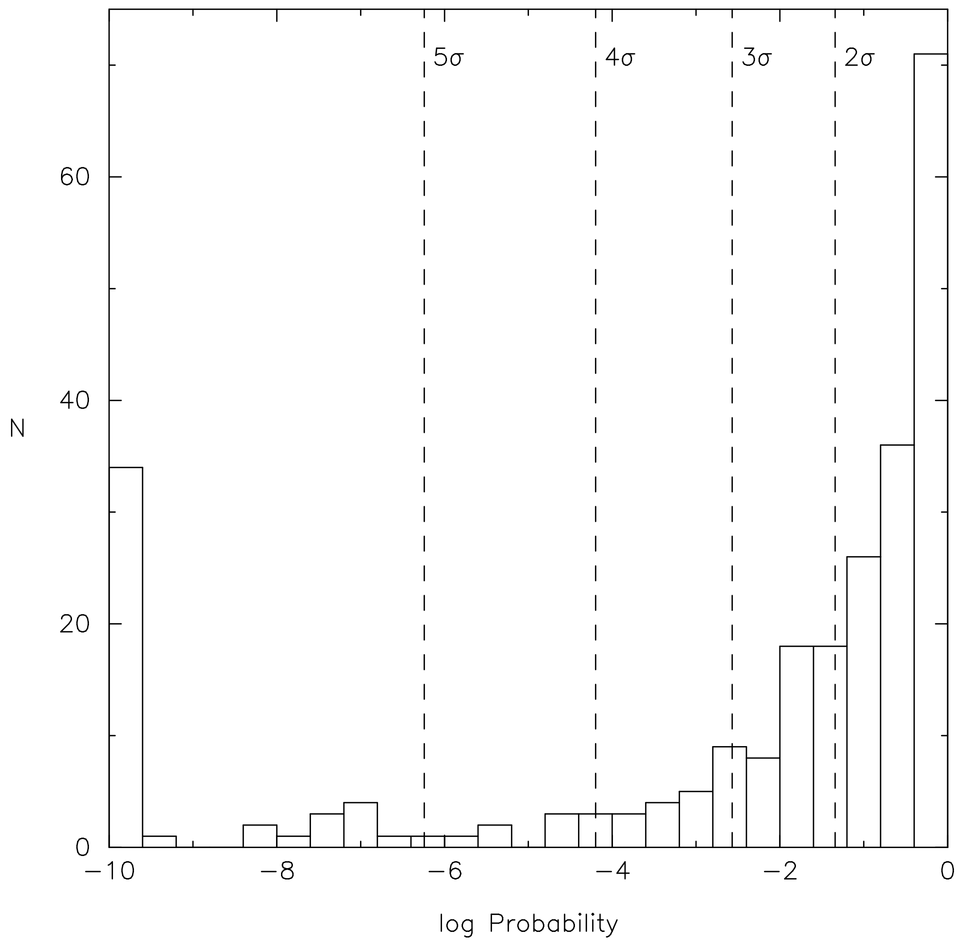

Fig. 5.— Histogram of the probabilities that the IRS low-resolution spectra between 12 and 35μm are better fitted by a Kurucz model atmosphere than a Kurcz model plus blackbody. Lower probabilities indicate the source is more likely to have an infrared excess from warm dust.

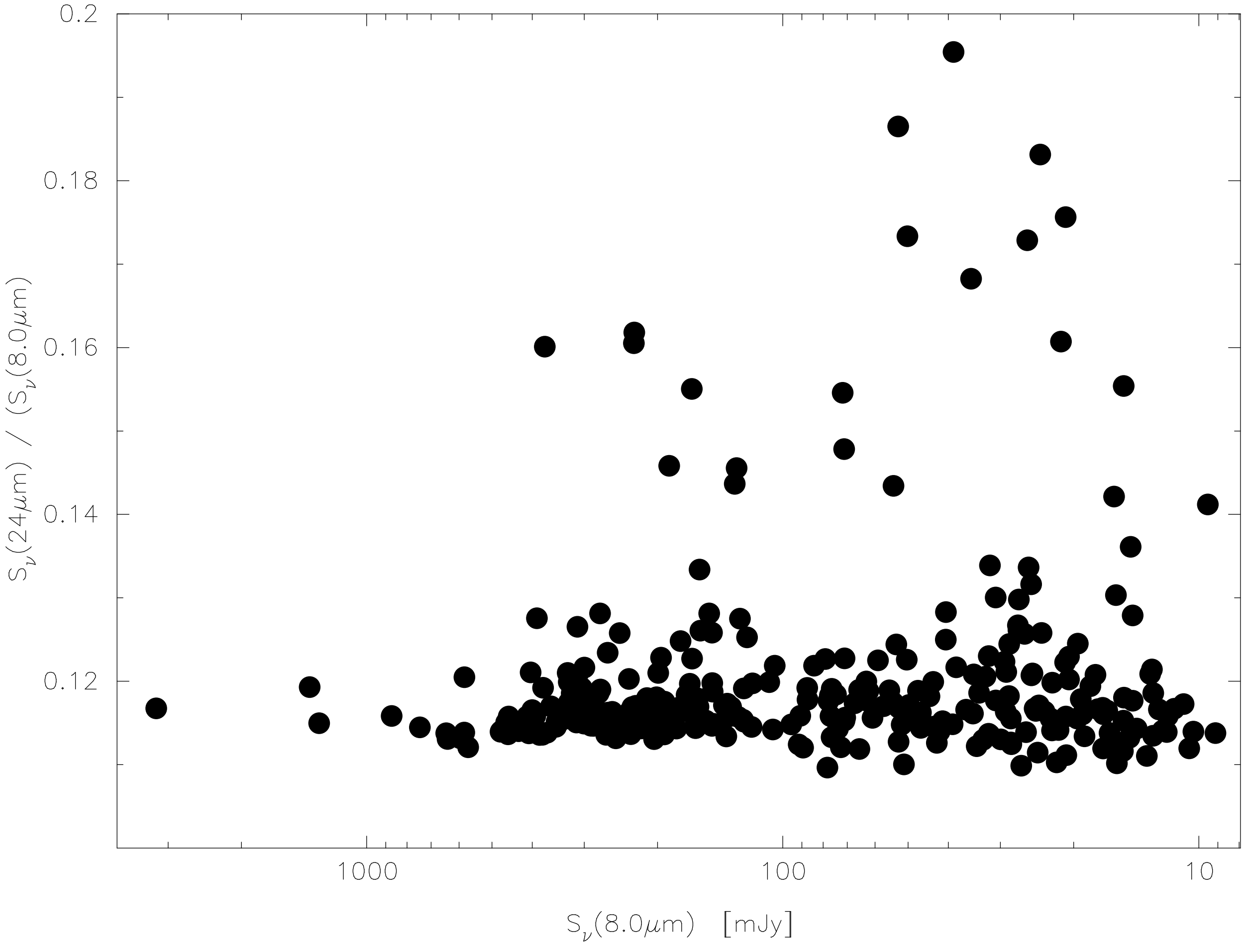

Fig. 6.— Color-magnitude diagram for the FEPS sources.

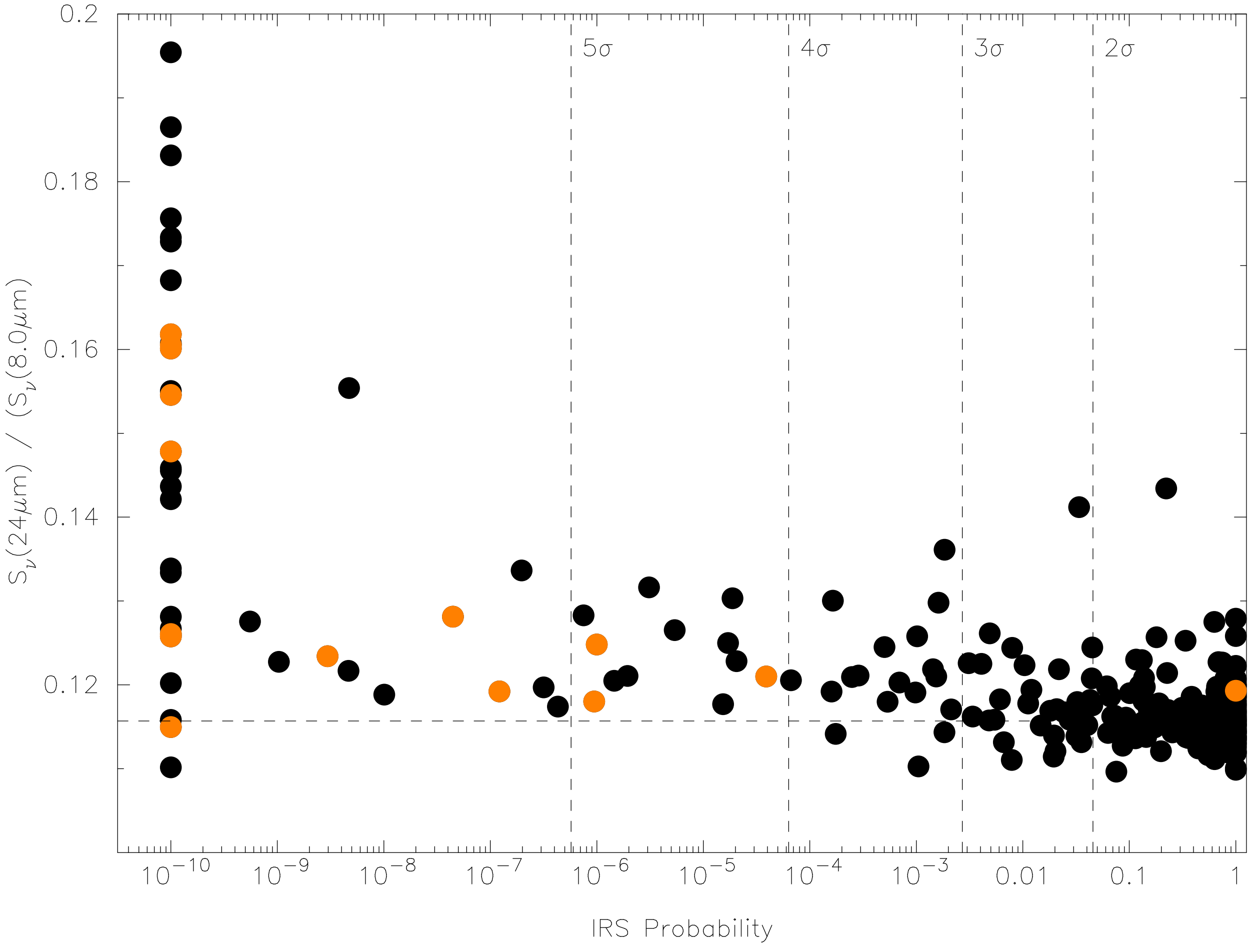

Fig. 7.— The ratio of the flux density in the MIPS 24μm band to the IRAC 8.0μm band as a function of the probability a Kurucz model can fit the observed IRS spectrum. The lower the probability, the more likely there is a infrared excess. Orange circles represent sources with a $\geq 3\sigma$ 70μm excess.

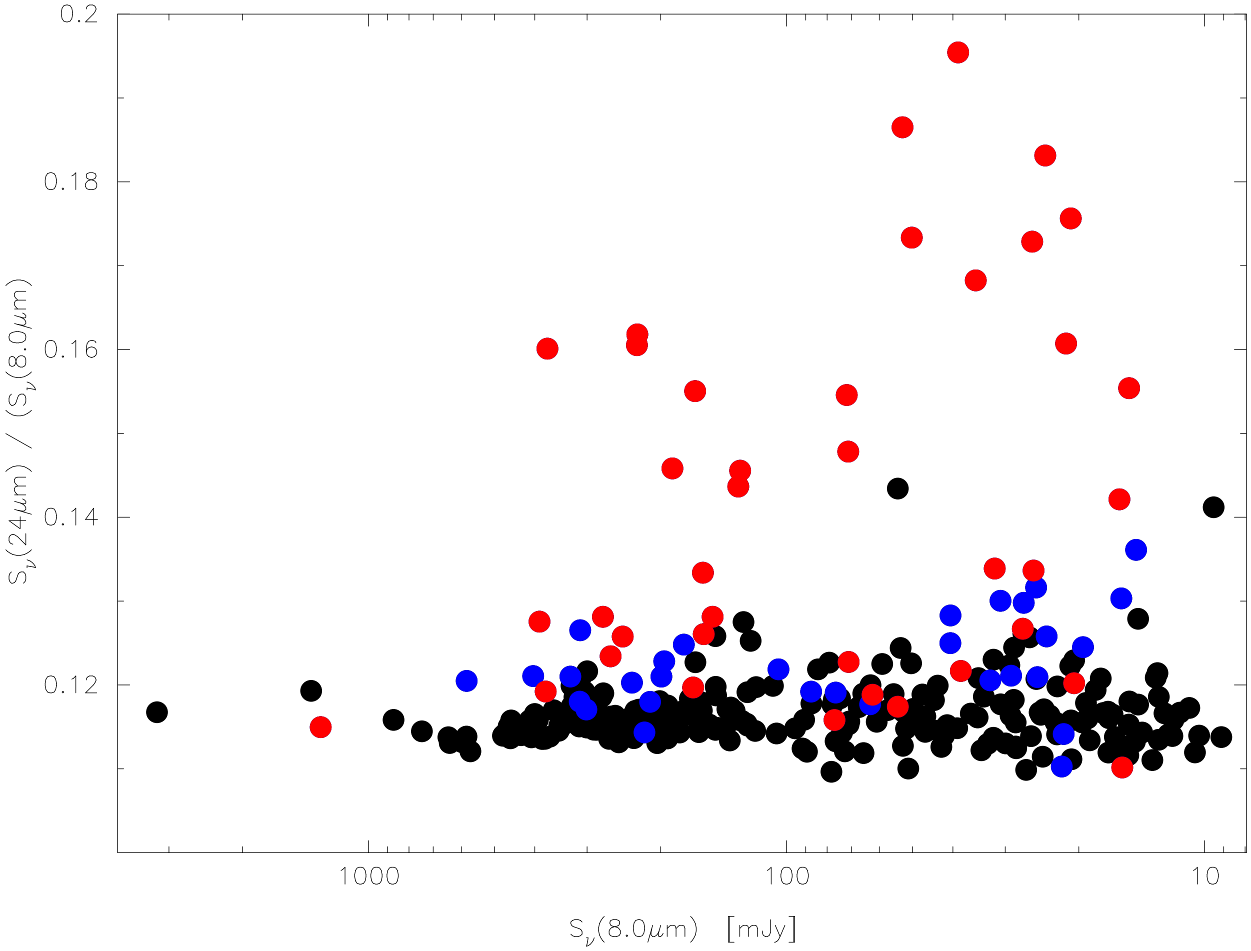

Fig. 8.— Color-magnitude diagram for the FEPS sources. Red circles are sources with a probablity of an IRS excess of $\geq 5\sigma$, and blue circles are for probabilities between 3σ and 5σ.

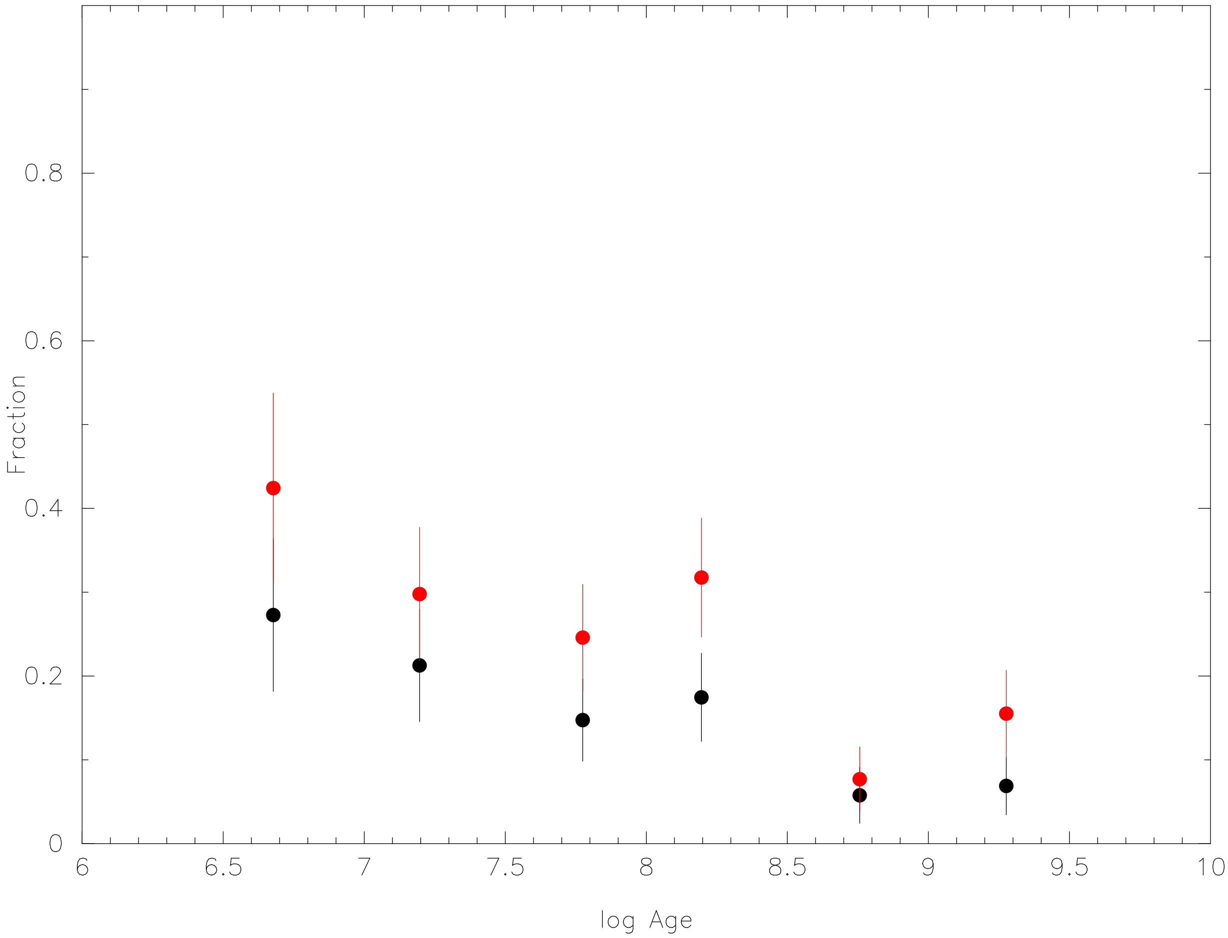

Fig. 9.— Fraction of sources with a IRS excess as a function of age. Red sources are for sources with a excess probability $\geq 5\sigma$, and blue sources are for $\geq 3\sigma$.

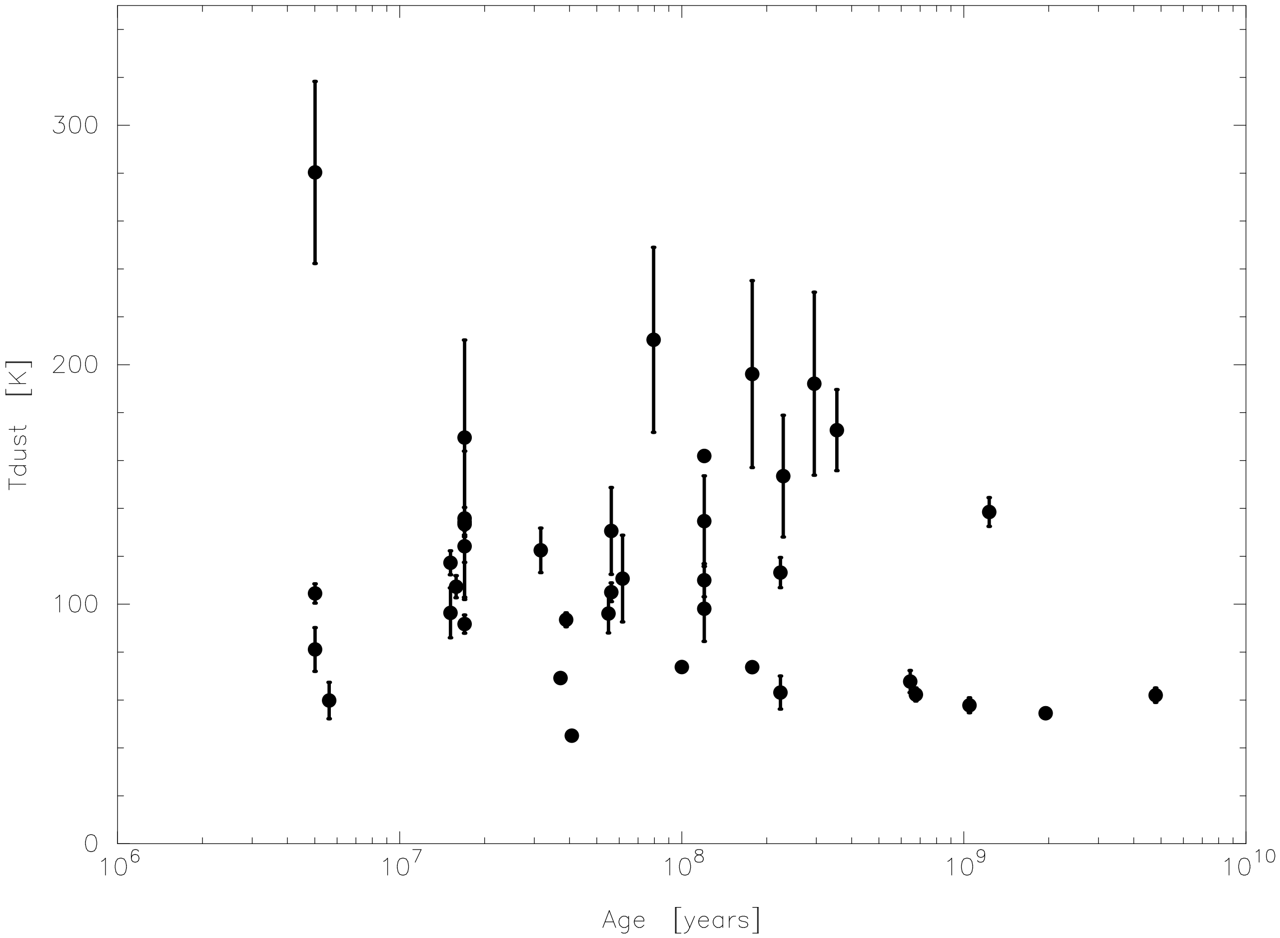

Fig. 10.— IRS-derived dust temperature as a function of stellar age for sources with $\geq 5\sigma$ probability of having an excess. The optically thick primordial disks have been removed.

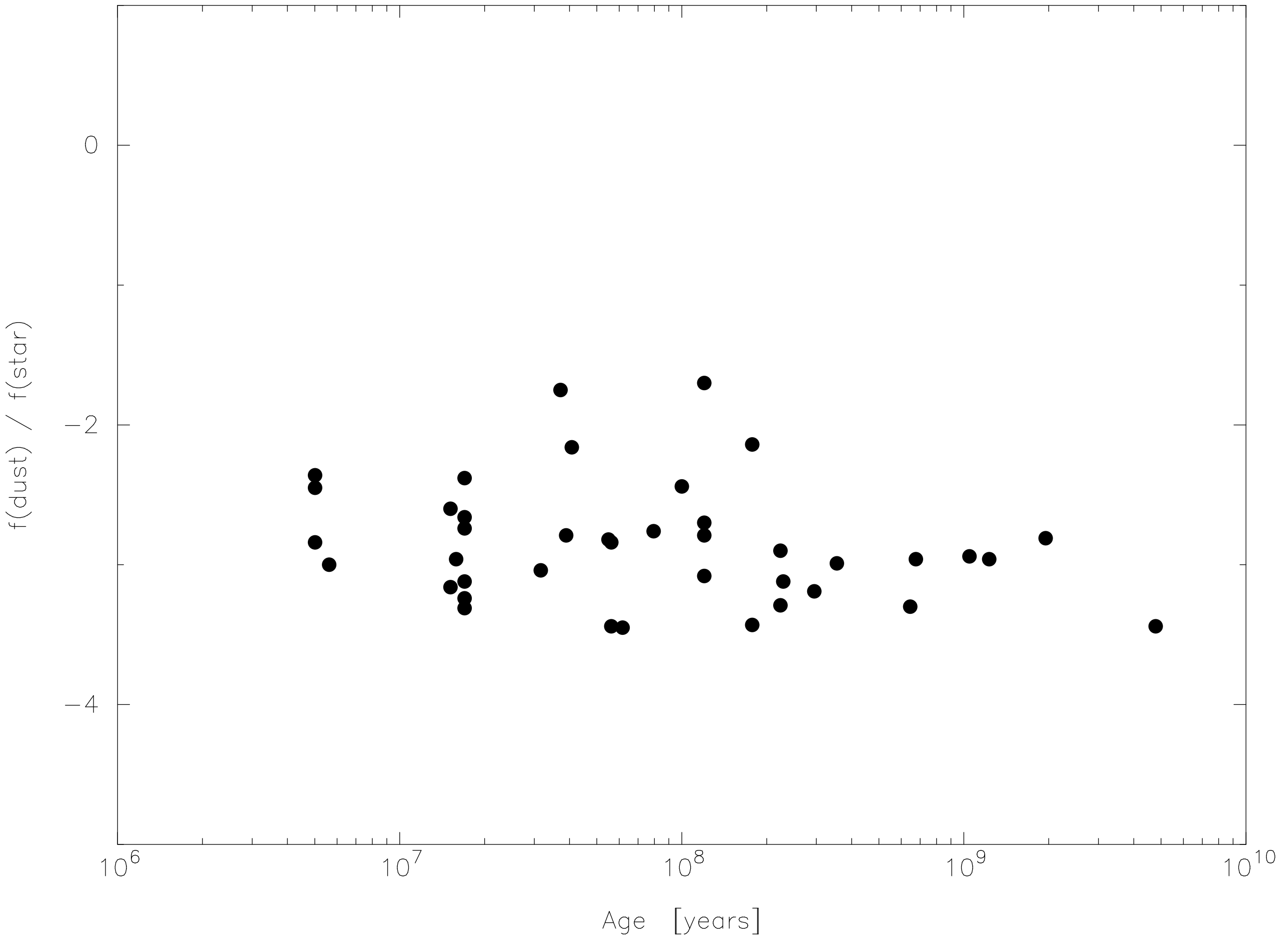

Fig. 11.— Fractional dust luminosity as a function of age.

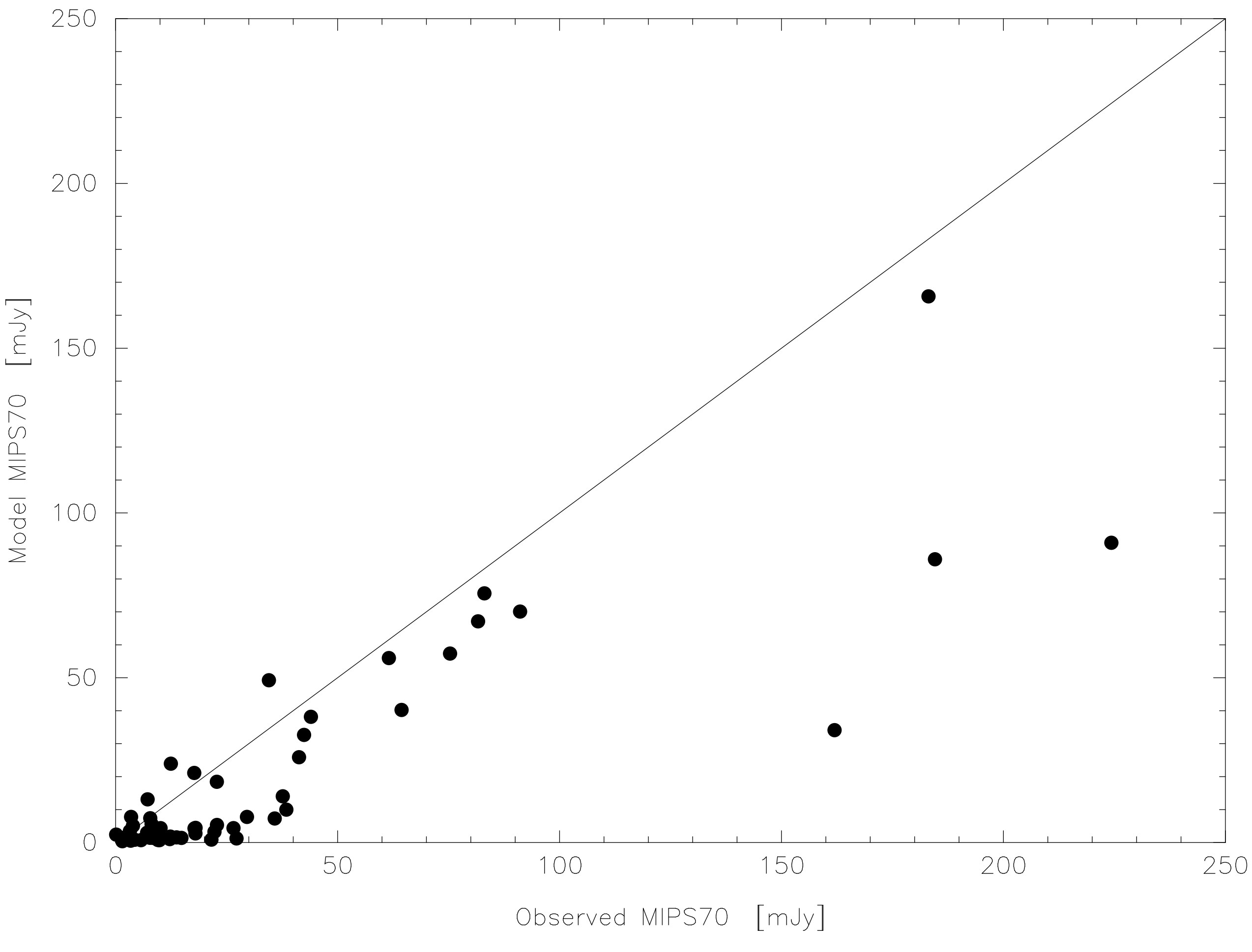

Fig. 12.— Model vs. observed 70μm fluxes. Model fluxes were derived from a blackbody fit to the IRS spectra.