CCD and wavelength, and, for our reductions, the inverse of the flat-field lamp spectrum (caused by the flat-field division).

Most of the response functions were determined using spline fits to a set of calibrated points at 50Å intervals (*cf.* Massey *et al.* 1988). Spectra reduced after about 1990 used the flux calibrated spectra (with points every ∼2Å) of standard stars available via Oke (1990). With these spectra, it was possible to simply divide the "true" spectrum by observed spectrum in order to produce an (inverse) response function. However, it was still necessary to extrapolate over stellar absorption features and to reduce the noise from the observed standard star spectra using a combination of polynomial fitting and smoothing.

A correction was also applied for the change in the atmospheric transmission versus wavelength as a function of airmass, since the object and standard star exposures were taken at different airmass.