

parameter. (2) Estimate the amount of scattered light filling in some BAL troughs. (3) Estimate accurate optical depths. (4) Determine the mechanism causing the BAL time variability. (5) Estimate (indirectly) the variability characteristics of the soft X-ray continuum. (6) Estimate relative abundances. (7) Estimate the physical conditions of the BALR as a function of outflow velocity. (8) Determine the relation between the BELR and BALR via variability in the BALs and BELs. (9) Confirm the identification of uncertain absorption features and the true distribution of column densities as a function of velocity, which are often confused by emission features. (10) Disentangle narrow absorption line systems and perhaps learn something about the nature of the narrow lines. (11) Estimate the structure and homogeneity of the BALR. (12) Set limits on the local covering factor and possibly the global covering factor by estimating the amount of emission originating in the BALR.

What residual intensity variations may not tell us is the distance between the BALR and the central source. This is because, unlike BEL variations, the BAL gas we observe is all along our line-of-sight (see §1.4).

Our data set of spectral and broadband monitoring can also potentially determine other quantities. From acceleration limits, *i.e.* changes in the velocity structure, we can set limits on the distance to the central source, and on the mechanism causing the acceleration. From the broadband data, we can learn more about the continuum variability of optically-selected QSOs, and determine if there are any differences between BALQSOs and non-BALQSOs.

Finally, although the spectroscopic observations were primarily focused on the BALs, we can use our data to estimate the character of BEL time variability in moderate-to-high redshift optically selected QSOs. This is a study which has been somewhat neglected, due to the small amplitude ($\Delta\text{magnitude} \sim 0.2$) of continuum variations in these objects, as well as their faint magnitudes, $V \gtrsim 17$, relative to lower redshift QSOs.