

## LIMITS ON THE POLARIZATION OF THE COSMIC MICROWAVE BACKGROUND RADIATION AT MULTIPOLES UP TO $l \sim 2000$

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### ABSTRACT

We report upper limits on the polarization of the cosmic microwave background radiation (CMBR) as measured with the Cosmic Background Imager (CBI), a 13 element interferometer that operates in the 26–36 GHz band and is sited on Llano de Chajnantor in northern Chile. The array consists of 90 cm Cassegrain antennas mounted on a steerable platform that can be rotated about the optical axis to facilitate polarization observations. The CBI employs single-mode circularly polarized receivers, and it samples multipoles from  $l \sim 400$  to  $l \sim 3500$ . The polarization data were calibrated on 3C 279 and Taurus A. The polarization observations consist of 278 hr of data on two fields taken in 2000, during the first CBI observing season. A joint likelihood analysis of the two fields yields three upper limits (95% confidence limit) for  $C_l^{\text{EE}} = C^{\text{EE}}l(l+1)/2\pi$  under the assumption that  $C_l^{\text{BB}} \equiv 0$ :  $49.0 \mu\text{K}^2$  ( $l = 603$ ),  $164 \mu\text{K}^2$  ( $l = 1144$ ), and  $630 \mu\text{K}^2$  ( $l = 2048$ ).

*Subject headings:* cosmic microwave background — cosmology: observations — galaxies: individual (Centaurus A) — supernovae: individual (Taurus A) — techniques: interferometric — techniques: polarimetric

### 1. INTRODUCTION

The cosmic microwave background radiation (CMBR) provides a unique means of testing many aspects of the standard model of the early universe. All variations of the standard model agree that the CMBR is the redshifted radiation from the initial plasma and that, as such, it contains clues to the fundamental characteristics of the universe (e.g., Kamionkowski & Kosowsky 1999; Hu & Dodelson 2002). This information resides in the spatial fluctuations of the total intensity and polarization of the CMBR. The past decade has seen the emergence of low-noise detector technologies that are propelling us into a new era of precision measurements of the characteristics of the CMBR. Observations have established the existence of intensity anisotropies with  $\delta T/T_0 \sim 10^{-5}$  on scales of  $\theta \sim 0.1\text{--}0.5^\circ$ , (e.g., Halverson et al. 2002; Netterfield et al. 2002; Kuo et al. 2004; Hinshaw et al. 2003; Readhead et al. 2004a). In contrast to the fluctuations in total intensity, polarization anisotropies are sufficiently small to have eluded detection until very recently (Kovac et al. 2002; Kogut et al. 2003; Leitch et al. 2005; Readhead et al. 2004b).

Standard models predict that Thomson scattering at the surface of the last scattering will polarize the fluctuations at the 10% level on scales of tens of arcminutes (e.g., Kamionkowski et al. 1997). The Cosmic Background Imager (CBI) is one of several experiments that have used the technique of radio interferometry to measure the spatial power spectrum of these fluctuations. Besides CBI and its sister instrument, the Degree Angular Scale Interferometer (DASI), recent experiments have employed a variety of methods to achieve sensitivities that approach cosmologically important levels: POLAR (Keating

et al. 2001), Saskatoon (Wollack et al. 1993; Netterfield et al. 1997), PIQUE (Hedman et al. 2001, 2002), CAPMAP (Barkats et al. 2005), and the Australia Telescope Compact Array (ATCA; Subrahmanyam et al. 2000). The CBI performed preliminary polarization observations in 2000, and the results of this work are reported here. This initial set of observations demonstrated the CBI's polarization capabilities (Cartwright 2003), and, on the basis of the success of this work, the telescope has been upgraded and dedicated to polarization observations since 2002 September; the detections that resulted from these observations are presented in a recent paper by Readhead et al. (2004b). Other ongoing polarization experiments include BOOMERANG (Masi et al. 2002), MAXIPOL (Johnson et al. 2003), COMPASS (Farese et al. 2004), QUEST (Piccirillo et al. 2002), *WMAP* (Bennett et al. 2003), and *Planck* (Delabrouille et al. 2002; Villa et al. 2002), the latter two of which are all-sky satellite missions.

### 2. THE COSMIC BACKGROUND IMAGER

The CBI is a 13 element interferometer that operates in the 26–36 GHz band (Padin et al. 2002). The array consists of 90 cm Cassegrain antennas mounted on a single, fully steerable platform. The antenna platform employs the standard altitude–azimuth axes, as well as a rotational degree of freedom about the telescope optical axis; this latter feature facilitates polarization observations. The platform allows a range of positions for the telescopes, permitting observations of anisotropies on multipoles  $l \sim 400\text{--}3500$ ; this range encompasses the scales over which standard models predict that much of the power in total intensity and polarization fluctuations is to be found. The observations reported in this paper concentrate on the  $400 < l < 2400$  region, to which the CBI is particularly well matched.

The CBI employs single-mode circularly polarized receivers. In these initial pioneering polarization observations with the CBI, the main focus was to determine the suitability of the instrument

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for polarization studies and to understand the primary sources of systematic error. To implement a polarization detection effort in parallel with the intensity observations that constituted the CBI's primary mission, we configured 12 receivers for left-circular polarization (LCP) and one receiver for right-circular polarization (RCP); the resulting array consisted of 66 total intensity (LL) and 12 cross-polarized (LR) baselines, all spanning  $400 < l < 3500$ . Each receiver has a quarter-wave plate that determines its polarization. The CBI was configured for polarization observations from 2000 January to 2000 October, at which point the single RCP antenna was converted back to LCP.

A single interferometer baseline measures a ‘‘visibility,’’ which is the Fourier transform of the intensity distribution on the sky. An LCP and RCP antenna pair forms an LR baseline, which measures the cross-polarized visibility  $\mathcal{V}^{\text{LR}}$  at a point  $\mathbf{u} = (u, v)$  in the aperture plane:

$$\mathcal{V}^{\text{LR}}(\mathbf{u}) = \int \int A(\mathbf{x} - \mathbf{x}_0)[Q(\mathbf{x}) - iU(\mathbf{x})]e^{2i\theta}e^{-2\pi i\mathbf{u} \cdot \mathbf{x}}d^2\mathbf{x}. \quad (1)$$

$A(\mathbf{x} - \mathbf{x}_0)$  is the primary beam pattern, which is assumed to be the same for both antennas that form the baseline and is centered at  $\mathbf{x}_0$  on the sky,  $\theta = \tan^{-1}(u/v)$ , and  $Q(\mathbf{x})$  and  $U(\mathbf{x})$  are Stokes parameters that describe the distribution of polarized flux on the sky. Although the integrals are evaluated over the entire sky,  $A(\mathbf{x} - \mathbf{x}_0)$  confines the signal of interest to the region of the sky in view of the primary beam. The interplay between the kernel of the Fourier transform and the primary beam determines the range of angular scales to which the baseline is sensitive; for an observation at wavelength  $\lambda$  on a baseline of length  $b$ , the synthesized beam  $\theta_s \sim |\mathbf{u}|^{-1} \sim \lambda/b$  determines the resolution, while the primary beam, for which  $\theta_p^{\text{FWHM}} = 46.5(\lambda/1 \text{ cm})$ , sets the field of view. Leitch et al. (2002) discuss the application of this method to polarization observations with DASI.

In the configuration for these observations, the CBI directly measures  $\mathcal{V}^{\text{LL}}(\mathbf{u})$  and  $\mathcal{V}^{\text{LR}}(\mathbf{u})$ .  $I(\mathbf{x})$  can be obtained from measurements of  $\mathcal{V}^{\text{LL}}(\mathbf{u})$  alone in the absence of circular polarization, but both  $\mathcal{V}^{\text{LR}}(\mathbf{u})$  and  $\mathcal{V}^{\text{RL}}(\mathbf{u})$  are required to obtain  $Q(\mathbf{x})$  and  $U(\mathbf{x})$ . Although an LR baseline does not directly measure RL, we can obtain it via rotation of the deck about the optical axis, since  $\mathcal{V}^{\text{LR}}(\mathbf{u}) = [\mathcal{V}^{\text{RL}}(-\mathbf{u})]^*$  (Conway & Kronberg 1969): a  $180^\circ$  deck rotation permits determination of  $Q(\mathbf{x})$  and  $U(\mathbf{x})$ . Although  $Q$  and  $U$  are necessary for imaging, for the likelihood analysis we use LR sampled over both halves of the  $(u, v)$  plane (§ 4).

The CBI antenna elements are Cassegrain dishes. Good polarization performance favors a clear aperture, but the Cassegrain optics do not impair the performance because the secondary of each telescope is supported by polystyrene feed legs that are transparent at centimeter wavelengths. Nonetheless, the optics do introduce contamination into the cross-polarized visibilities. To reduce cross talk, the antennas are surrounded by cylindrical shields; unpolarized blackbody emission from the ground is polarized upon scattering from the insides of the antenna shields to the antenna feeds. This spurious spillover signal dominates the cross-polarized visibilities on short baselines, but the lead-trail observing technique eliminates this contamination (§ 3). To test for the presence of spurious off-axis polarization, we measured the instrumental polarization at the half-power points of the primary beam in the four cardinal directions. A  $\chi^2$  test demonstrated that the instrumental polarization at these points is consistent with that at the antenna boresight, showing that the instrumental polarization properties do not degrade rapidly as one moves off-axis.

## 2.1. Polarization Calibration

The cross-polarized visibilities must be calibrated to measure the gain and instrumental polarization. To first order, the raw cross-polarized visibility for the baseline using antennas  $j$  and  $k$  is given by

$$\mathcal{V}_{jk}^{\text{LR}}(\mathbf{u}) = G_{jk} \{ [\tilde{Q}(\mathbf{u}) - i\tilde{U}(\mathbf{u})]e^{-2i\psi} + \epsilon_{jk}\tilde{I}(\mathbf{u}) \}, \quad (2)$$

where  $G_{jk}$  and  $\epsilon_{jk}$  denote the *baseline*-based instrumental gain and polarization, respectively, and  $\tilde{I}(\mathbf{u})$ ,  $\tilde{Q}(\mathbf{u})$ , and  $\tilde{U}(\mathbf{u})$  are the Fourier transforms of  $I(\mathbf{x})$ ,  $Q(\mathbf{x})$ , and  $U(\mathbf{x})$ . The terms  $G_{jk}$  and  $\epsilon_{jk}$  are both complex quantities and must be evaluated for each of the 10 CBI bands. The instrumental polarization, or ‘‘leakage,’’ permits the total intensity to contaminate the cross-polarized visibilities; for observations of the CMBR, uncorrected leakage will cause measurements of polarization fluctuations at a particular  $l$  to be contaminated by the total intensity fluctuations at the same  $l$ . Here  $\psi$  denotes the deck orientation about the optical axis. To determine  $G_{jk}$  and  $\epsilon_{jk}$ , we observe a source of known polarization and total intensity at a variety of deck orientations, and the change in  $\psi$  modulates the first term of equation (2) relative to the second. Observations at a minimum of two different deck orientations are required to obtain both  $G_{jk}$  and  $\epsilon_{jk}$ .

To calibrate the LR visibilities for the CMBR deep fields, equation (2) is evaluated at each  $(u, v)$  point, using the values of  $G_{jk}$  and  $\epsilon_{jk}$  determined from the calibration observations, together with measurements of  $\tilde{I}(\mathbf{u})$ , to isolate  $\tilde{Q}(\mathbf{u}) - i\tilde{U}(\mathbf{u})$ . During the 2000 observing season, the array configuration and observational strategy precluded LL matches for all LR visibilities; for short baselines ( $b \sim 100 \text{ cm}$  or  $l \sim 600$ ), where we expect the greatest signal, the loss of data was not substantial, but for several of the longer baselines ( $b \sim 300 \text{ cm}$  or  $l \sim 1900$ ), the lack of LL matches precluded the use of all data.

In the present observations, the amplitude of the instrumental polarization averages  $\sim 8\%$  for all baselines and all channels and can approach 20%. The receiver was modeled to understand the source of the leakage. The instrumental polarization is dominated by bandpass errors in the quarter-wave plates; at the edges of the  $\sim 30\%$  fractional band, the insertion phase of the quarter-wave plates departs from  $\lambda/4$  by several percent. In addition, assembly errors can cause the plate orientation to depart from the ideal  $45^\circ$  with respect to the rectangular guide that follows it. A receiver model that incorporates these errors shows excellent agreement with the measured leakage (Fig. 1). The modeled receiver characteristics that give rise to  $\epsilon$  are stable, so we infer that the measured leakage is stable as well. High signal-to-noise ratio (S/N) measurements of the leakage at regular intervals demonstrated that it remained stable over timescales spanning many months.

The polarization data were calibrated with observations of extragalactic radio source 3C 279 and the supernova remnant Taurus A (the Crab Nebula). 3C 279, a bright extragalactic radio source, served as the primary polarization calibrator. With  $I \sim 25 \text{ Jy}$  and  $m = |P|/I \sim 10\%$  at 31 GHz [where  $|P| = (Q^2 + U^2)^{1/2}$ ] and no significant extended emission on CBI scales, 3C 279 permits quick calibration observations. 3C 279 is variable, however, so it was monitored at monthly intervals throughout the polarization campaign with the NRAO Very Large Array<sup>3</sup> (VLA) at 22.46 and 43.34 GHz. 3C 279 showed some activity

<sup>3</sup> The National Radio Astronomy Observatory is a facility of the National Science Foundation operated under cooperative agreement by Associated Universities, Inc.

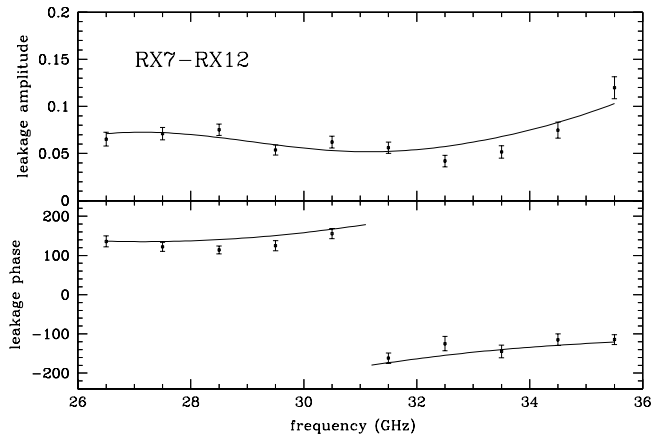


FIG. 1.—Comparison of leakage model fit to leakage data for the baseline between RX7 and RX12, across all 10 channels. The top panel shows the leakage amplitude  $\epsilon_4$ , in units for which 1.0 corresponds to 100% leakage, while the bottom panel shows the leakage phase  $\epsilon_\phi$ . Points represent measurements of the leakage, while lines show the model.

during the January–May period; at 22.46 GHz, its fractional polarization changed by  $\sim 20\%$  ( $\delta m \sim 0.02$ ), while its position angle rotated by  $10^\circ$  during the same period. These changes were approximately linear at 22.46 GHz and occasionally discontinuous at 43.34 GHz. Although the VLA observations yield  $I$ ,  $Q$ , and  $U$ , we transferred only the fractional polarization  $m$  and the position angle  $\chi$  [where  $2\chi = \tan^{-1}(U/Q)$ ] to the CBI. This choice permits us to use daily measurements of the total intensity of 3C 279 with the CBI to set the flux density scale for the polarization observations. The absolute uncertainty of the total intensity calibration is 4%.

We required two interpolations to apply the VLA values for  $m$  and  $\chi$  to the CBI observations. The first interpolation transfers  $m$  and  $\chi$  from the two VLA channels to the 26–36 GHz band. Measurements of the total intensity of 3C 279 in the 10 CBI bands show that the total intensity is well characterized by a power law, and in light of this uniform behavior we made a simple linear fit to both  $m$  and  $\chi$ . The statistical uncertainty on this fit is typically  $<5\%$  in amplitude and  $<3^\circ$  in position angle. The second interpolation transfers  $m$  and  $\chi$  from the dates of the VLA observations to the intervening CBI observations. Again, a linear interpolation was used. While the statistical uncertainty in this interpolation is generally small ( $<5\%$  in amplitude and  $<3^\circ$  in position angle) and trivial to compute, the systematic uncertainty is harder to estimate, particularly for the 43.34 GHz data; the changes in  $m$  and  $\chi$  between VLA observations at 43.34 GHz in one interval are quite high ( $\sim 20\%$  and  $\sim 10^\circ$ ), although only  $\sim 1\%$  of the CMBR data were taken during this interval. However, regular measurements of the total intensity of 3C 279 with the CBI show that it does not undergo excursions beyond those seen in the VLA data, so we assume that the temporal variations in the polarization characteristics do not exceed those in the VLA data. The measurement uncertainties in the VLA data are typically 3%, so the uncertainties in the interpolated VLA data can be as high as 8%.

3C 279 was observed nearly every night at a pair of deck positions separated by  $90^\circ$ ; each observation lasted 5 minutes and was accompanied by a trailing field to measure contamination from ground spillover (§ 3). The total uncertainty in the 3C 279 calibration is typically 9%, of which 8% arises from the uncertainty in the VLA data, 3% results from the uncertainty in the CBI LR observations, and 4% arises from the flux scale, which is set by the uncertainty in the CBI’s LL calibration.

TABLE 1  
GAUSSIAN MODEL COMPONENTS FOR TAURUS A AT 31 GHz

Component	$S_\nu$ (Jy)	$x_0^a$ (arcsec)	$y_0^a$ (arcsec)	$\sigma^b$ (arcmin)	$b/a^b$	$\phi^b$ (deg)
$I$ .....	355.3	0.0	0.0	3.58	0.66	–50
$Q$ .....	14.9	–48.8	116.9	2.93	0	83
$U$ .....	–23.9	–30.1	128.2	2.28	0.52	56

<sup>a</sup>  $x_0$  and  $y_0$  are positions of the centroids of the model components, measured with respect to that for the total intensity.

<sup>b</sup>  $\sigma$ ,  $b/a$ , and  $\phi$  are the major axis width, axial ratio, and orientation, respectively, of the elliptical Gaussian model to which each component was fit.

Taurus A served as the polarization calibrator for nearly 40% of the polarization data. Taurus A is marginally resolved by the CBI, so we required a simple model for its morphology. There are no published data on Taurus A’s polarized emission at 31 GHz, so we transferred the calibration on 3C 279—obtained directly from a nearly contemporaneous VLA observation—to the Taurus A observations and derived a model. Our Taurus A model consists of single elliptical Gaussian components for each of  $I$ ,  $Q$ , and  $U$ ; these model components are shown in Table 1, and this simple model is applicable over ranges of  $|u| \sim 100$ –500 and the 26–36 GHz band. The spectral indices for the two polarized components were constrained to be that of the total intensity:  $\alpha = -0.3$ , where  $S_\nu \propto \nu^\alpha$ .

We performed a number of supporting observations to assess the accuracy of the polarization calibration. We included 3C 273 in the VLA monitoring campaign, and observations of 3C 273 with the CBI provide a test of the internal consistency of the polarization calibration. 3C 273 is a  $\sim 25$  Jy,  $\sim 5\%$  polarized source at centimeter wavelengths, and the polarization we recover from CBI observations of 3C 273 is consistent with the VLA observations within the statistical and systematic uncertainties of the calibration. Cross-checks on observations of 3C 279 provide estimates of the uncertainty on the calibration with the Taurus A model in Table 1; using Taurus A as a calibrator, we recover  $m$  and  $\chi$  for 3C 279 to within  $\sim 10\%$  and  $\sim 5^\circ$ , respectively. We regard these values as the uncertainties on the polarization calibration.

### 3. OBSERVATIONS OF THE COSMIC MICROWAVE BACKGROUND RADIATION

The data presented here were obtained from deep observations of two fields, the 08<sup>h</sup> field ( $\alpha = 08^{\text{h}}44^{\text{m}}40^{\text{s}}$ ,  $\delta = -3^\circ 10' 00''$ ) and the 20<sup>h</sup> field ( $\alpha = 20^{\text{h}}48^{\text{m}}40^{\text{s}}$ ,  $\delta = -3^\circ 30' 00''$ ); measurements of total intensity fluctuations in these fields have been reported by Mason et al. (2003). These fields are a subset of a group of four fields spaced at equal intervals in right ascension<sup>4</sup> that were selected for minimum contamination from diffuse Galactic emission. Both fields are at Galactic latitudes  $>24^\circ$ . Each field is the size of a single beam, or  $45'$  FWHM at the band center (§ 2). Simple extrapolations from Haslam’s 408 MHz maps suggested that for both fields, the polarization fluctuations from synchrotron emission at 1 cm on CBI scales would be smaller than the CMBR polarization fluctuations (Haslam et al. 1982).

The observations presented in this paper were obtained between 2000 January and 2000 October. The 08<sup>h</sup> field was observed from January through the end of May, and the 20<sup>h</sup> field was observed from August through the end of October, at which point the array was dedicated to total intensity observations until

<sup>4</sup> The CBI’s elevation limit of  $43^\circ$  constrains the time on-source to  $\sim 6$  hr day<sup>–1</sup>.

2002 September. This work encompassed 99 nights of observations: 44 nights on the 08<sup>h</sup> field and 55 nights on the 20<sup>h</sup> field, which yielded 130 hr and 148 hr of data, respectively.

The 08<sup>h</sup> field observations spanned two array configurations, while the 20<sup>h</sup> field was observed with a third. The weather at the Chajnantor site was generally excellent when observations were not precluded by snowstorms, and less than 1% of the data were flagged.

The observing strategy was guided by several considerations. The visibilities measured on the short baselines are contaminated by ground spillover, so we observed fields in pairs separated by 8<sup>m</sup> in right ascension and differenced the pairs offline to reject the common spillover contribution<sup>5</sup> (Mason et al. 2003). To within the uncertainties of the visibilities, the LL and LR visibilities show no evidence of spillover after differencing. The observations were performed at night and when the moon was >60° from the fields. Each lead-trail pair was tracked in constant parallactic angle, and, after each pair, the deck position was advanced by either 20° or 30°. Each 8<sup>m</sup> scan consists of ~50 8.4 s integrations; ~15% of each scan is lost to calibrations and slews.

We performed a number of consistency tests on the CMBR data prior to the likelihood analysis. We first applied a jackknife test to assess the accuracy of the noise estimates. The visibility uncertainty for each 8<sup>m</sup> scan was estimated from the scatter in the ~50 integrations in the scan. For the jackknife test, the real and imaginary visibilities were sorted into two interleaved sets corresponding to alternating dates, and at each  $(u, v)$  point each set was averaged over time. The two sets were then differenced on a point-by-point basis in the  $(u, v)$  plane. We are most concerned about effects on the shortest baselines, for which we expect the greatest signal and, conversely, for which the spillover contamination is greatest. To that end, for the 08<sup>h</sup> field we computed  $\chi^2_\nu = 1.03$  with  $\nu = 590$  (probability to exceed [PTE] = 30%) for the real components and  $\chi^2_\nu = 1.11$  (PTE = 3%) for the imaginary ones. Similarly, for the 20<sup>h</sup> field visibilities ( $\nu = 720$ ), we find  $\chi^2_\nu = 1.00$  (real) and  $\chi^2_\nu = 0.97$  (imaginary), both of which are consistent with unity.

We were also concerned that systematic errors in the polarization of the calibrator sources—particularly the Taurus A model—would give rise to errors in the calibration of the CMBR data. Since  $\sigma_V \propto |G_{jk}|$ , we used the visibility uncertainties as a proxy for the amplitude component of the LR gain calibration; we averaged the LR visibility uncertainties for the CMBR data and compared them to those for LL, the gain calibration of which we believe to be accurate to ~4%. After accounting for a slight (4%) excess in system noise for RX12—the orthogonally polarized antenna that is common to all LR visibilities—we find that  $\langle\sigma_{LR}\rangle \sim \langle\sigma_{LL}\rangle$  to within 10%, which is consistent with the results of the calibration cross-checks discussed in § 2.1.

#### 4. ADDITIONAL OBSERVATIONS OF POLARIZED SOURCES

We observed several polarized sources to assess the polarization performance of the CBI. Centaurus A (NGC 5128) is a nearby active galaxy that exhibits a rich variety of polarized structure over a range of angular scales at centimeter wavelengths. W44 has several janskies of polarized emission at 1 cm, and its size of ~30' is comparable to the primary beam of the CBI. We discuss these examples here.

We observed Centaurus A for 6.8 hr with the CBI. Figure 2 shows the CBI map of Centaurus A's double inner lobes, along

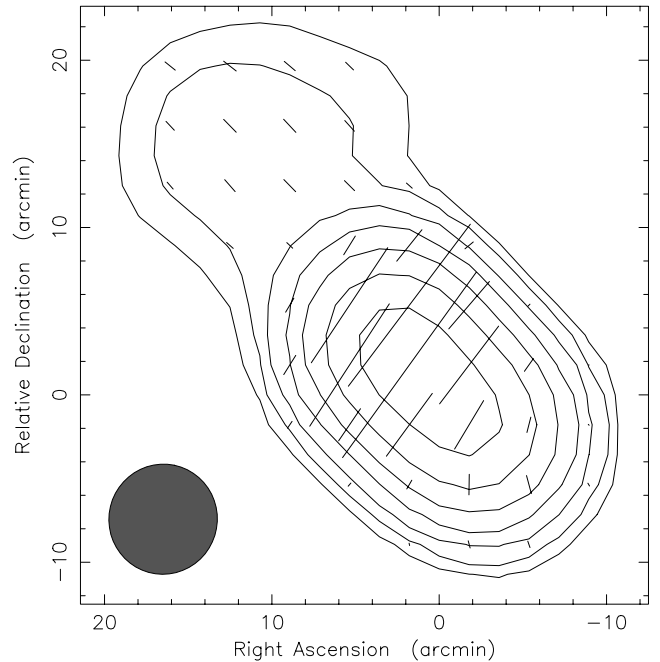


FIG. 2.—Clean LL CBI map of the double inner lobes of Centaurus A at 31 GHz from 2000 June 8, centered at  $\alpha = 13^{\text{h}}35^{\text{m}}27^{\text{s}}.600$ ,  $\delta = -43^{\circ}01'08''.800$ . Contours are shown for total intensity (1, 2, 4, 8, 16, 32, and 64 percent), while the polarization magnitude and position angle are represented by lines. The map peak is 23.6 mJy beam<sup>-1</sup>, and the beam FWHM is 6''.59 × 6''.48 at -11'. The southern edge of the northern middle lobe appears as the dim feature at the upper left, while the bright oval-shaped region of emission encompasses the northern and southern inner lobes. The text discusses the distribution of polarized flux.

with the southernmost edge of the northern middle lobe. The image is centered on the northern end of the double inner lobe, at which point the total intensity peaks at 20.1 Jy, the fractional polarization reaches 12%, and the position angle is -36°.

While the total intensity of the southern lobe resembles that of the northern lobe—it peaks at 18.7 Jy beam<sup>-1</sup>—the polarization characteristics of the southern lobe are strikingly different: the fractional polarization reaches 3.6% at the total intensity peak, at which point the position angle is approximately -37°. Junkes et al. (1993) present observations of the inner lobes of Centaurus A at 6.3 GHz with the Parkes 64 m telescope; the authors report that at the northern inner lobe the fractional polarization peaks at 13%, while at the peak of the total intensity of the southern inner lobe the polarization rises to only ~5% at the southernmost edge of the lobe. The position angle across the two inner lobes is -70° <  $\chi$  < -33°, and it wraps around to ~+5° along the southern slope of the southern inner lobe. The CBI results are consistent with these findings.

We observed W44 for 2.6 hr with the CBI. Figure 3 shows the CBI map of W44 after having been restored with a 8.6 × 7' beam. The remnant has a pear-shaped shell, with a distinct asymmetry arising from the steep density gradient in the immediate neighborhood of the remnant (Cox et al. 1999). The CBI maps show that the fractional polarization peaks at ~33% on the northwestern slope of the source, and across the center of the source it is relatively uniform at 10%–12%. While the position angle varies across the source, it is roughly uniform at ~60° across most of the emission in total intensity. Kundu & Velusamy used the NRAO 140 feet telescope to map W44 at 10.7 GHz with a 3' beam (Kundu & Velusamy 1972). The authors report that the fractional polarization peaks at ~20% along the northeast edge, and it remains uniform over the dominant region of emission along the

<sup>5</sup> The positions given above are those of the leading fields.

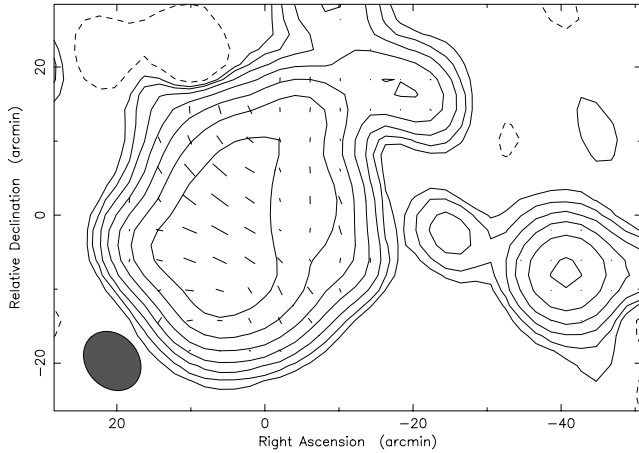


FIG. 3.—Clean LL CBI map at 31 GHz from 2000 June 23 of supernova remnant W44 (on the left) and the Galactic H II region G34.3+0.1 (on the right), centered at  $\alpha = 18^{\text{h}}56^{\text{m}}02^{\text{s}}270$ ,  $\delta = +01^{\circ}21'57''.229$ . Contours are shown at  $-1, 1, 2, 4, 8, 16, 32$ , and  $64$  percent intensity. The map peak is  $4.23 \text{ Jy beam}^{-1}$ , and the beam FWHM is  $8''.64 \times 7''.03$  at  $41^{\circ}$ . The H II region is unpolarized, and the lack of polarization  $>0.5\%$  in the CBI map of G34.3+0.1 demonstrates that the CBI does not create spurious polarization greater than this level.

east side of the source. At the peak of emission in total intensity, the authors find that the fractional polarization  $m \sim 20\%$ . The neighborhood of W44 contains a Galactic H II region that provides a key test of the CBI's polarization capabilities. The emission from this source, G34.3+0.1, is due to free-free emission, so the source should be unpolarized. The fractional polarization at the total intensity of the emission is  $\sim 0.5\%$ , so we conclude that the CBI is not creating spurious polarization at greater than this level. These tests gave us great confidence in the potential of the CBI for polarization observations, and they were an important factor in our decision to upgrade the instrument to carry out a focused program of polarization observations.

### 5. LIKELIHOOD ANALYSIS OF THE POLARIZATION DATA

The method of maximum likelihood was used to test the data for the presence of a hypothetical signal. The “likelihood” of the data  $\mathbf{x}$  given a theory  $\mathbf{q}$  is given by

$$\mathcal{L}(\mathbf{x}|\mathbf{q}) = \frac{1}{\pi^{N_d} |\mathbf{C}(\mathbf{q})|} \exp[-\mathbf{x}^t \mathbf{C}^{-1}(\mathbf{q}) \mathbf{x}],$$

where  $\mathbf{x}$  is a data vector of length  $N_d$  and the covariance matrix  $\mathbf{C}$  quantifies the correlations between these data for the model under test. In this analysis,  $\mathbf{x}$  is a vector that contains the real and imaginary components of the LR visibilities  $\mathcal{V}^{\text{LR}}(\mathbf{u})$  that populate both halves of the  $(u, v)$  plane.  $\mathbf{C}(\mathbf{q})$  consists of a theoretical correlation  $\mathbf{M}$  and a diagonal noise matrix  $\mathbf{N}$ :  $\mathbf{C}(\mathbf{q}) = \mathbf{M}(\mathbf{q}) + \mathbf{N}$ . The model  $\mathbf{q}$  that maximizes the likelihood or, equivalently, the  $\log$  likelihood

$$\ln \mathcal{L}(\mathbf{x}|\mathbf{q}) = -N_d \ln(\pi) - \ln[|\mathbf{C}(\mathbf{q})|] - \mathbf{x}^t \mathbf{C}(\mathbf{q}) \mathbf{x}$$

is regarded as the model that is most consistent with the data. The model may be a function of  $l$ ; we assume a model with *flat band power*, for which  $q_i = C_l^{\text{EE}} = l(l+1)/2\pi C_l^{\text{EE}}$  is constant for the  $l$  range defined for band  $i$ . Several authors (e.g., Hobson et al. 1995; Myers et al. 2003) discuss techniques for applying the method of maximum likelihood to visibility data; we have implemented aspects of the approaches discussed by these authors with the assumption that  $C_l^{\text{BB}} \equiv 0$ .

The deep field observations described in this work yielded  $\sim 10^7$  visibilities, each of which corresponds to an 8.4 s integration for a single baseline and channel, so the visibilities were averaged to reduce the covariance matrix to tractable proportions. The visibilities were averaged in three passes. First, the  $\sim 50$  8.4 s integrations in each  $8^{\text{m}}$  scan were averaged to form a single visibility for that scan. The uncertainty for the scan-averaged visibility was computed from the scatter in the constituent visibilities. This procedure introduces a downward bias in the noise, so the elements of the noise covariance matrix  $\mathbf{N}$  were scaled upward by 6% to correct for this bias (Mason et al. 2003). Next, all the visibilities for all nights were averaged by  $(u, v)$  point, and finally, to truncate the size of the covariance matrix, the visibilities were averaged by band. The band average has the potential to bias the best-fit band powers, so we analyzed two simulated sets of low S/N data, one with a single  $\Delta\nu = 1$  GHz band centered at the middle of the CBI band and another with the entire  $\Delta\nu = 10$  GHz band averaged to a single  $\Delta\nu = 1$  GHz band that was centered at the same frequency as the first set.

We found that the upper limits obtained from the two sets of data were consistent; this should be the case, as the data are dominated by noise. The final data set consisted of 185 and 149 discrete  $(u, v)$  points for the  $08^{\text{h}}$  and  $20^{\text{h}}$  fields, respectively. To expedite the likelihood calculation, these visibilities were sorted into three bins based on  $|\mathbf{u}|$ ; because of the spacing between the antennas, this binning scheme resulted in one based on physical length: band 1 incorporated the 100 and 104 cm baselines, band 2 contained the 173 and 200 cm baselines, and band 3 contained the remaining long baselines. The resulting upper limits do not correct for correlations between these bands.

Simulations provide insights about the effects of errors in the calibration, so we simulated data with errors in the complex gain  $G_{jk}$  and complex leakage  $\epsilon_{jk}$  (eq. [2]). The simulations demonstrated that substantial errors in the gain phase  $G_\phi$  ( $\delta G_\phi \sim 6^\circ$ , or 10% of a radian) result in negligible changes ( $\sim 1\%$ ) to the best-fit band power, while changes to the gain amplitude  $G_A$  scale the best-fit band powers quadratically, as expected. Systematic errors in the leakage calibration are of particular concern because they can mimic real polarization in the CMBR. These simulations show that the errors in the leakage phase  $\epsilon_\phi$  of the instrumental polarization do not affect the best-fit band powers (for fixed nonzero  $\epsilon_A$ ), while errors in the leakage amplitude  $\epsilon_A$  tend to *increase* the best-fit band powers regardless of whether they overestimate or underestimate the true leakage amplitude; this must be the case, since the power in fluctuations is purely additive. Errors in  $\epsilon_A$  contribute in quadrature with the intrinsic polarization on the sky:  $C_l^{\text{EE}} \rightarrow C_l^{\text{EE}} + \delta\epsilon_A^2 C_l^{\text{TT}}$ . A 20% error in the amplitude of 10% instrumental polarization, for example, tends to bias the amplitude of the best-fit band power  $(C_l^{\text{EE}})^{1/2}$  *upward* by  $<2\%$  for a generic standard cosmology. We are therefore confident that the band powers reported in this work are not contaminated by errors in the leakage correction by more than this level.

Since we report upper limits in this paper, our primary concern is that systematic calibration errors do not cause us to *underestimate* these limits. The simulations demonstrated that of the four types of calibration errors ( $G_A$ ,  $G_\phi$ ,  $\epsilon_A$ , and  $\epsilon_\phi$ ), only a systematic error in the gain amplitude can bias the limits downward, and, as noted previously, a variety of cross-checks demonstrated that the error on  $G_A$  is 10%. All the sources of uncertainty—the assumptions for the likelihood calculation and errors in the instrumental polarization calibration—tend to result in overestimates of the best-fit band powers; we are confident that the

limits reported herein do not underestimate the sky signal beyond the uncertainty in the gain calibration.

## 6. RESULTS

The 278 hr of deep field data yielded several upper limits on  $C_l^{EE}$ . Table 2 lists the 95% confidence limit results for the measurements of the two fields and the joint fit to the fields; these were obtained by integrating the likelihood from  $q \equiv 0$ . We have assumed that  $C_l^{BB} \equiv 0$ . For each band, the band center is the peak of the summed diagonal elements of  $\mathbf{M}$ , while the band width is the FWHM of the summed diagonal elements of  $\mathbf{M}$ . As a cross-check, the likelihood routine was modified to address  $C_l^{TT}$ ; it was tested on the short-baseline 08<sup>h</sup> field data, for which it yielded  $(C_l^{TT})^{1/2} = 66.8_{-11.1}^{+14.1}$   $\mu\text{K}$ . This value is consistent with the value obtained by the CBI for nearly the same  $l$  range:  $(C_l^{TT})^{1/2} = 62.9_{-7.9}^{+11.3}$   $\mu\text{K}$  (Padin et al. 2001); the two sets of data have differing  $(u, v)$  coverage, so the two measurements are not identical.

These upper limits are consistent with the rapidly burgeoning body of CMBR polarization data. The limit at  $l = 603$  is consistent with limits in the same region from DASI and CBI. The limits for the higher  $l$  bins are consistent with predictions for  $C^{EE}$  for generally accepted families of models. The limits also provide constraints on polarized emission from galactic synchrotron emission and polarized point sources in these regions and on these scales. This pioneering polarization effort with the CBI

TABLE 2  
UPPER LIMITS ON  $C_l^{EE}$ , 95% CONFIDENCE

Band	$l_{\min}$	$l_c$	$l_{\max}$	08 <sup>h</sup> $\sqrt{q}$ ( $\mu\text{K}$ )	20 <sup>h</sup> $\sqrt{q}$ ( $\mu\text{K}$ )	Joint $\sqrt{q}$ ( $\mu\text{K}$ )
1.....	446	603	779	14.1	8.1	7.0
2.....	930	1144	1395	21.2	15.9	12.8
3.....	1539	2048	2702	45.3	27.7	25.1

provided great confidence in the polarization capabilities of the instrument, and it was a central consideration in our decision to upgrade the CBI for a dedicated polarization program.

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